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## HESS J1943+213: a candidate extreme BL Lacertae object

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### ABSTRACT

**Context.** The H.E.S.S. Cherenkov telescope array has been surveying the Galactic plane for new VHE (>100 GeV) gamma-ray sources.

**Aims.** We report on a newly detected point-like source, HESS J1943+213. This source coincides with an unidentified hard X-ray source IGR J19443+2117, which was proposed to have radio and infrared counterparts.

**Methods.** We combine new H.E.S.S., *Fermi*/LAT and Nançay Radio Telescope observations with pre-existing non-simultaneous multi-wavelength observations of IGR J19443+2117 and discuss the likely source associations as well as the interpretation as an active galactic nucleus, a gamma-ray binary or a pulsar wind nebula.

**Results.** HESS J1943+213 is detected at the significance level of  $7.9\sigma$  (post-trials) at RA(J2000) =  $19^{\text{h}}43^{\text{m}}55^{\text{s}} \pm 1^{\text{s}}_{\text{stat}} \pm 1^{\text{s}}_{\text{sys}}$ , Dec(J2000) =  $+21^{\circ}18'8'' \pm 17''_{\text{stat}} \pm 20''_{\text{sys}}$ . The source has a soft spectrum with photon index  $\Gamma = 3.1 \pm 0.3_{\text{stat}} \pm 0.2_{\text{sys}}$  and a flux above 470 GeV of  $(1.3 \pm 0.2_{\text{stat}} \pm 0.3_{\text{sys}}) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$ . There is no *Fermi*/LAT counterpart down to a flux limit of  $6 \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$  in the 0.1–100 GeV energy range (95% confidence upper limit calculated for an assumed power-law model with a photon index  $\Gamma = 2.0$ ). The data from radio to VHE gamma-rays do not show any significant variability.

**Conclusions.** The lack of a massive stellar counterpart disfavors the binary hypothesis, while the soft VHE spectrum would be very unusual in case of a pulsar wind nebula. In addition, the distance estimates for Galactic counterparts places them outside of the Milky Way. All available observations favor an interpretation as an extreme, high-frequency peaked BL Lac object with a redshift  $z > 0.14$ . This would be the first time a blazar is detected serendipitously from ground-based VHE observations, and the first VHE AGN detected in the Galactic Plane.

**Key words.** galaxies: active – BL Lacertae objects: individual: IGR J19443+2117 – gamma rays: general

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## 1. Introduction

The High Energy Stereoscopic System (H.E.S.S.) is an array of four imaging atmospheric Cherenkov telescopes situated in the Khomas Highland of Namibia (Aharonian et al. 2006b). H.E.S.S. collaboration has been conducting a systematic scan of the Galactic Plane, increasing the survey sensitivity and expanding the catalog of very high energy (VHE,  $>100$  GeV) sources with new intrinsically faint or distant objects. In the survey, which now covers most of the Galactic Plane as seen from the Southern Hemisphere, a rich population of VHE galactic sources has been discovered (Aharonian et al. 2005c, 2006c; Chaves 2009).

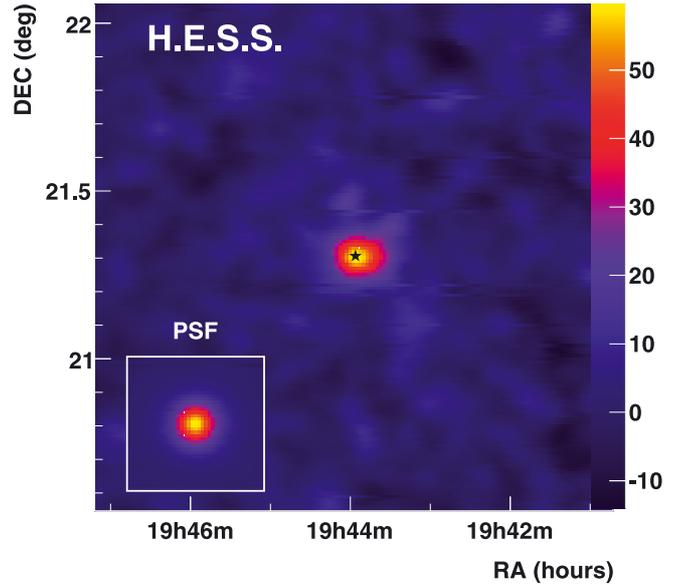
The majority of known VHE sources are extended beyond the  $\sim 0.1^\circ$  H.E.S.S. point spread function (PSF). The few sources in the Galactic Plane that appear point-like are often associated with VHE gamma-ray emitting high-mass X-ray binaries (HMXB) which include the very well established binaries PSR B1259-63 (Aharonian et al. 2005b), LS 5039 (Aharonian et al. 2006d) and LS I +61 303 (Albert et al. 2006). The point-like VHE source HESS J0632+057 is now a strong candidate for a HMXB system following a recent multi-wavelength campaign (Aharonian et al. 2007; Hinton et al. 2009). In addition several young pulsar wind nebulae (PWN), including the Crab Nebula (Aharonian et al. 2006b), are also unresolved by H.E.S.S.. Outside of the Galactic Plane the point-like sources are associated with active galactic nuclei (AGN). More than 30 AGN, mostly BL Lacertae (BL Lac) objects, have been detected in VHE gamma-rays up to now<sup>1</sup>. None have been detected yet serendipitously in the Galactic Plane survey.

Here, we present the discovery of a new unresolved VHE gamma-ray source, HESS J1943+213, located in the Galactic Plane and spatially coincident with an unidentified INTEGRAL source, IGR J19443+2117 (Bird et al. 2007). This paper is organized as follows. In Sect. 2, we present the H.E.S.S. data analysis and results. In Sect. 3, we present the sources found within the H.E.S.S. source error circle and evaluate the plausibility of each of them being a counterpart of the H.E.S.S. source. This section also includes an analysis of archival INTEGRAL and *Fermi* data, as well as of data taken recently by the Nançay Radio Telescope. Assuming that all these sources are counterparts to the H.E.S.S. source, Sect. 4 discusses the possible association of HESS J1943+213 with an HMXB, PWN and AGN. The conclusions are set out in Sect. 5.

## 2. H.E.S.S. data

The source was initially discovered in the Galactic plane scan data collected between 2005 and 2008. Dedicated observations were performed between May and August 2009. These observations were taken in *wobble* mode, where the telescopes point in a direction typically at an offset of  $0.5^\circ$  from the nominal target position. The offset angle of the scan-mode observations varies between  $0.5^\circ$  and  $2^\circ$ . After applying selection cuts to the data, to reject periods affected by poor weather conditions and hardware problems, the total observation time used for analysis amounts to 38.4 h. This is equivalent to live-time of a 24.8 h on axis due to reduced acceptance for runs (individual continuous observation segments of  $\sim 28$  min) taken with large offsets from the nominal source position. The mean zenith angle of the observations is  $49^\circ$ .

<sup>1</sup> Based on the online catalog for TeV Astronomy TeVCat located at <http://tevcat.uchicago.edu>, provided by Scott Wakely & Deirdre Horan.



**Fig. 1.** Excess map of the field around the position of HESS J1943+213, smoothed by the PSF of the instrument (68% containment radius of  $0.064$  deg).

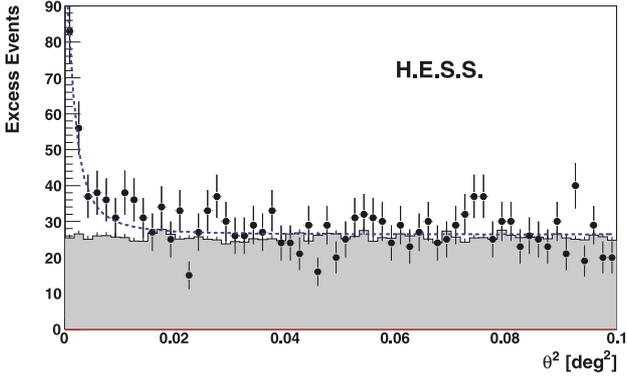
The data were calibrated according to Aharonian et al. (2004). Energies are reconstructed taking the effective optical efficiency into account (Aharonian et al. 2006b). Event reconstruction and separation of gamma-ray-like events from cosmic-ray background were made using the *Model Analysis* and standard cuts (de Naurois & Rolland 2009), in which shower images of all triggered telescopes are compared to a pre-calculated model by means of a log-likelihood minimization.

For the spectral analysis, on-source data are taken from a circular region of radius  $\theta^2 = 0.01$  deg<sup>2</sup> around the source position, and the background is subtracted using the event background rate estimated with the reflected background model (Aharonian et al. 2006b) for which several off regions at the same distance to the center of the camera as the target position are used (excluding the region close to the source).

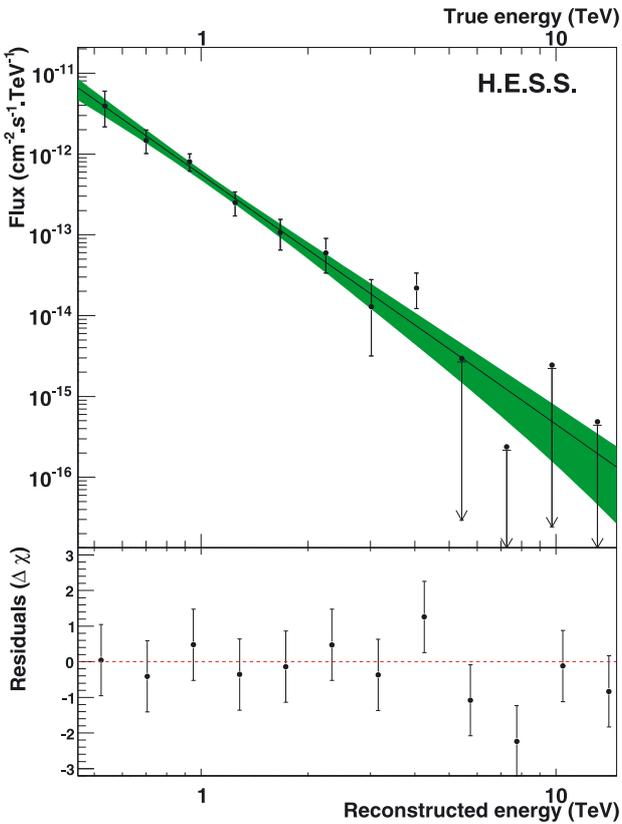
A total of  $N_{\text{On}} = 281$  on-source events and  $N_{\text{Off}} = 4086$  off-source events are measured. The on-off normalization factor is  $\alpha = \Omega_{\text{On}}/\Omega_{\text{Off}} = 0.0379$ , where  $\Omega$  is the solid angle of the respective on- and off-source regions. The observed excess is  $N_\gamma = N_{\text{On}} - \alpha N_{\text{Off}} = 126$   $\gamma$ -rays, corresponding to a significance of  $8.9\sigma$  pre-trials standard deviations according to Eq. (6) from Li & Ma (1983).

In Fig. 1 we plot an excess map of the sky around the HESS J1943+213 position smoothed by the PSF of the instrument. The angular distribution of events shown in Fig. 2 is consistent with that of a point source. After correcting for trials (Aharonian et al. 2006c) the statistical significance of the detection is  $7.9\sigma$ . A fit of a point-like source convolved with the H.E.S.S. PSF of the excess yields a position at the 68% confidence level,  $\text{RA}(J2000) = 19^{\text{h}}43^{\text{m}}55^{\text{s}} \pm 1^{\text{s}}_{\text{stat}} \pm 1^{\text{s}}_{\text{sys}}$ ,  $\text{Dec}(J2000) = +21^{\circ}18'8'' \pm 17''_{\text{stat}} \pm 20''_{\text{sys}}$ . The source is located in the Galactic plane at  $l = 57.76^\circ$  and  $b = -1.29^\circ$ . The intrinsic source extension, calculated at the best fit position is smaller than  $2.8'$  at  $3\sigma$  confidence level.

Figure 3 shows the time-averaged differential spectrum which was derived using the forward folding technique with an energy threshold of 470 GeV. Between 470 GeV and  $\sim 6$  TeV this spectrum is well described (Log-likelihood chance probability 45%) by a power-law  $dN/dE = \Phi_0(E/1 \text{ TeV})^{-\Gamma}$  with photon

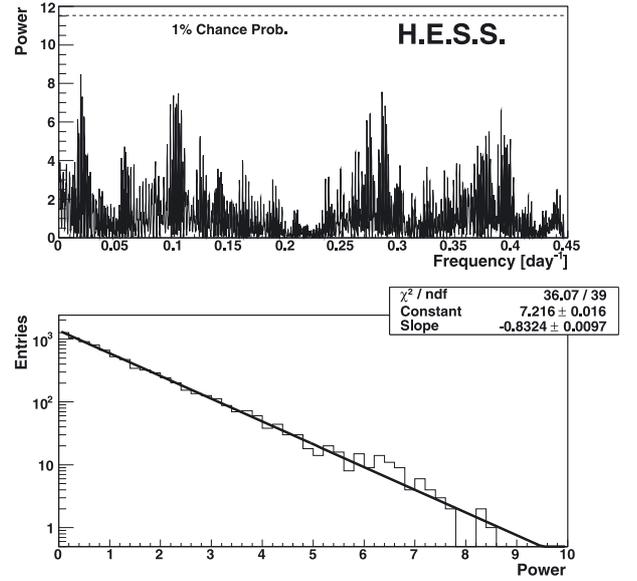


**Fig. 2.** Distribution of squared angular distance ( $\theta^2$ ) for gamma-like events and normalized background events. The dashed line denotes a point source profile.

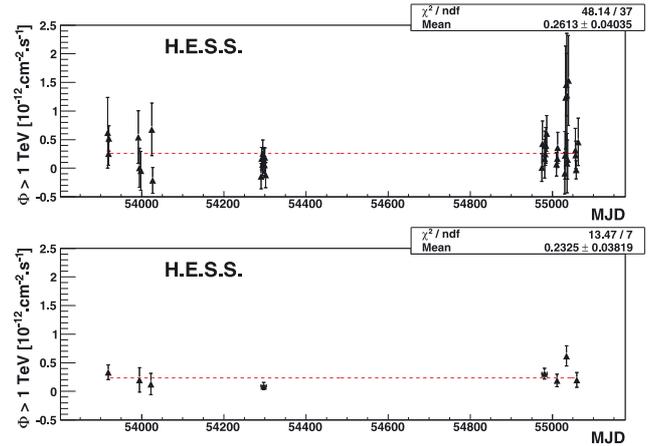


**Fig. 3.** Time-averaged VHE spectrum observed from HESS J1943+213. The shaded area represents the  $1\sigma$  confidence level of the fitted spectrum using a power-law hypothesis. Only the statistical errors are shown and the upper limits are 68% confidence level. The lower panel shows the residual of the spectrum fit, which are the difference, in each reconstructed energy bin, between expected and observed number of excess events, normalized to uncertainty on the latter ( $(N_{\text{obs}} - N_{\text{exp}})/\Delta N_{\text{exp}}$ ).

index  $\Gamma = 3.1 \pm 0.3_{\text{stat}} \pm 0.2_{\text{sys}}$  and normalization at 1 TeV  $\Phi_0(1 \text{ TeV}) = (5.6 \pm 0.8_{\text{stat}} \pm 1.1_{\text{sys}}) \times 10^{-13} \text{ cm}^{-2} \text{ s}^{-1} \text{ TeV}^{-1}$ . The 68% confidence level upper limits for the four highest-energy bins shown in Fig. 3 were calculated using the method of [Feldman & Cousins \(1998\)](#). The integrated flux above 470 GeV is  $I(>470 \text{ GeV}) = (1.3 \pm 0.2_{\text{stat}} \pm 0.3_{\text{sys}}) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$ , which corresponds to  $\sim 2\%$  of the Crab Nebula flux above the same energy threshold ([Aharonian et al. 2006b](#)). In Fig. 4 we show a



**Fig. 4.** *Top:* Lomb-Scargle periodogram of HESS J1943+213 based on the run-by-run lightcurve. The dashed line corresponds to the 1% chance probability level. *Bottom:* distribution of power, adjusted by an exponential distribution. In the case of a pure white noise, a slope of  $-1$  is expected.



**Fig. 5.** Integral flux  $I(>1 \text{ TeV})$  of HESS J1943+217, measured by H.E.S.S. during each observing night (*top*) and averaged per month (*bottom*). Only the statistical errors are shown. The dashed line corresponds in each case to the best fit to a constant level.

Lomb-Scargle periodogram for HESS J1943+213 and in Fig. 5 lightcurves. There is no significant variability in the data with chance probabilities of the null hypothesis of steady emission respectively of 10.4% and 6.2% for the night by night and month by month lightcurve (using statistical uncertainties only).

All the H.E.S.S. results presented in this paper have been cross-checked with an independent calibration procedure and a different background discrimination method ([Ohm et al. 2009](#)).

### 3. Counterparts

We searched for possible H.E.S.S. source counterparts in a number of radio, IR, optical and X-ray catalogs. Table 1 contains the position and spectral information with uncertainties for each source discussed in this section by reason of finding them

**Table 1.** The multi-wavelength fluxes and flux upper limits of potential counterparts of HESS J1943+213.

Observatory	Band	RA (J2000) [hh mm ss.s]	Dec (J2000) [° ' "]	Position Error 90%	$\theta$ <sup>(9)</sup> ["]	$\alpha = \Gamma - 1$	Flux [erg cm <sup>-2</sup> s <sup>-1</sup> ]	Ref.
NVSS <sup>(1)</sup>	1.4 GHz	19 43 56.14	+21 18 26.9	0.45", 0.56"	24.7	0.32 ± 0.02 <sup>(10)</sup>	(1.44 ± 0.05) × 10 <sup>-15</sup> <sup>(12)</sup>	a, b
NRT <sup>(1)</sup>	1.4 GHz	19 43 56	+21 18 26	15" <sup>(8)</sup>	22.8	0.6 ± 0.2	(1.55 ± 0.02) × 10 <sup>-15</sup> <sup>(12)</sup>	
2MASS <sup>(2)</sup>	2.2 μm (K)	19 43 56.244	+21 18 23.38	<0.2"	23.2	–	2.37 × 10 <sup>-12</sup> <sup>(12,13)</sup>	c
Swift/UVOT	0.54 μm (V)	19 43 56.20	+21 18 22.95	–	22.5	–	<1.6 × 10 <sup>-13</sup> <sup>(12,13)</sup>	g
ROSAT/HRI <sup>(3)</sup>	0.1–2 keV	19 43 56.2	+21 18 24	6"	23.2	– <sup>(14)</sup>	(3.82 ± 0.55) × 10 <sup>-11</sup> <sup>(14)</sup>	d
<i>Chandra</i> <sup>(4)</sup>	0.3–10 keV	19 43 56.23	+21 18 23.6	0.64"	23.2	0.83 ± 0.11	(2.9 <sup>+2.4</sup> <sub>-0.5</sub> ) × 10 <sup>-11</sup>	e
Swift/XRT <sup>(5)</sup>	2–10 keV	19 43 56.20	+21 18 22.95	3.53"	22.5	1.04 ± 0.12	(1.83 ± 0.04) × 10 <sup>-11</sup>	f, g
Swift/BAT <sup>(6)</sup>	15–150 keV	19 43 54.35	+21 18 08.2	2.88"	9.1	– <sup>(11)</sup>	(1.8 ± 0.2) × 10 <sup>-11</sup>	h
Swift/BAT <sup>(5)</sup>	14–195 keV	19 43 55.2	+21 18 07.2	–	2.9	1.1 <sup>+0.3</sup> <sub>-0.2</sub>	(2.9 <sup>+0.5</sup> <sub>-0.5</sub> ) × 10 <sup>-11</sup>	k
INTEGRAL/IBIS <sup>(7)</sup>	20–100 keV	19 44 09.36	+21 18 25.2	4.4"	201	1.12 ± 0.22	(1.12 ± 0.22) × 10 <sup>-11</sup>	g, i
INTEGRAL/SPI <sup>(7)</sup>	100–200 keV	19 44 23.13	+21 16 36.4	–	404	–	<2.6 × 10 <sup>-11</sup>	j
<i>Fermi</i> /LAT	0.1–100 GeV	19 43 56.23	+21 18 23.6	–	23.2	1.0 <sup>(15)</sup>	<7 × 10 <sup>-12</sup>	
<i>Fermi</i> /LAT						0.5 <sup>(15)</sup>	<1 × 10 <sup>-12</sup>	
H.E.S.S.	>470 GeV	19 43 55	+21 18 8	34" <sub>stat</sub> , 21" <sub>sys</sub>	0	2.1 ± 0.3 <sub>stat</sub> ± 0.2 <sub>sys</sub>	(1.8 <sup>+0.6<sub>stat</sub>+0.5<sub>sys</sub></sup> <sub>-0.4<sub>stat</sub>-0.4<sub>sys</sub></sub> ) × 10 <sup>-12</sup>	

**Notes.** (1) NVSS J194356+211826; (2) 2MASS J19435624+2118233; (3) 1RXH J194356.2+211824; (4) CXOU J194356.2+211823; (5) SWIFT J1943.5+2120; (6) PBC J1943.9+2118; (7) IGR J19443+2117; (8) Error for right ascension only; (9) angular distance from the centroid of H.E.S.S. source position; (10) spectral index given for all combined radio measurements between 327 and 4850 MHz, listed in SPECFIND V2.0 (Vollmer 2009); (11) Hardness ratio Rate(30–150 keV)/Rate(14–30 keV) = 1.0 ± 0.4; (12)  $EF(E)$ ; (13) the flux is not corrected for extinction; (14) unabsorbed flux obtained from ROSAT/HRI count-rate using a Mission Count Rate Simulator PIMMS, assuming  $N_{\text{H}}$  and  $\alpha_{\text{X}}$  identical with *Chandra*. The flux error includes uncertainties of  $N_{\text{H}}$  and  $\alpha_{\text{X}}$ ; (15) the spectral index was assumed for upper limit calculation.

**References.** (a) Condon et al. (1998); (b) Vollmer (2009); (c) Skrutskie et al. (2006); (d) ROSAT Scientific Team (2000); (e) Tomsick et al. (2009); (f) Malizia et al. (2007); (g) Landi et al. (2009); (h) Cusumano et al. (2010); (i) Bird et al. (2010); (j) Bouchet et al. (2008); (k) Baumgartner et al. (2011).

plausible counterpart of HESS J1943+213. Figure 6 shows the H.E.S.S. source position confidence contours along with the candidate counterpart positions.

### 3.1. X-ray counterpart

The H.E.S.S. source is located within the error circle (4.4') of an unidentified hard X-ray INTEGRAL/IBIS source IGR J19443+2117, the centroid position of which is about 3.3 arcmin away from the H.E.S.S. position. A counterpart to IGR J19443+2117 was detected in the soft X-ray domain by *Chandra* (CXOU J194356.2+211823; Tomsick et al. 2009) and Swift (SWIFT J1943.5+2120; Landi et al. 2009) with exposure times of 4.8 ks and 11 ks, respectively. The same X-ray source was detected in the past by ROSAT/HRI (1RXH J194356.2+211824, exposure 1.3 ks, ROSAT Scientific Team 2000)

The combined power-law fit to Swift/XRT and INTEGRAL/IBIS data gives a spectral index  $\alpha_{\text{X}} = 1.04$ <sup>(2)</sup> and fluxes  $F_{2-10 \text{ keV}} = 1.83 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ ,  $F_{20-100 \text{ keV}} = 1.12 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  (Landi et al. 2009). Within the error of the cross-calibration constant between XRT and IBIS equal 0.6, the IBIS data is in agreement with that of XRT.

The *Chandra* observation (performed two years later in 2008), is consistent with the Swift results and gives  $F_{0.3-10 \text{ keV}} = 2.9 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  with a spectral index  $\alpha_{\text{X}} = 0.83$  (Tomsick et al. 2009). The oldest soft X-ray observation of the source comes from ROSAT/HRI. The Mission Count Rate

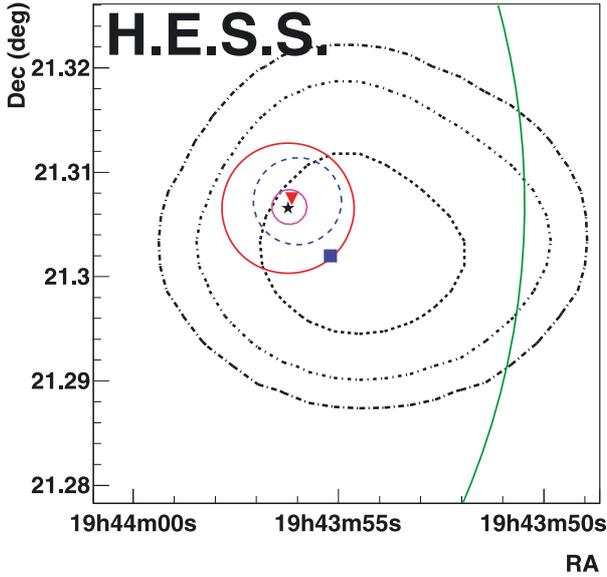
Simulator PIMMS was used in order to obtain the flux from the HRI count-rate. Assuming a hydrogen column density (see Sect. 3.2) and spectral index identical to the *Chandra* source, the flux equals  $F_{0.1-2 \text{ keV}} = 3.82 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ , this is in agreement with both *Chandra* and Swift fluxes within errors.

The INTEGRAL/SPI observations provide an upper limit on the source flux above 100 keV of  $F_{>100 \text{ keV}} < 2.6 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  (Bouchet et al. 2008). An X-ray source consistent with all above X-ray observations is also present in the Swift/BAT hard X-ray Catalog (PBC J1943.9+2118, Cusumano et al. 2010) with  $F_{14-150 \text{ keV}} = 1.8 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ , and in the 58 months Swift/BAT catalog (Baumgartner et al. 2011) with  $F_{14-195 \text{ keV}} = 2.86 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  and  $\alpha_{\text{X}} = 1.08$ . A power-law model fit to the spectrum provided by Baumgartner et al. (2011)<sup>3</sup> yields  $\chi^2/\text{d.o.f.} = 2.4/6$ , whereas a cut-off power-law model fit yields  $\chi^2/\text{d.o.f.} = 3/5$ . There is no evidence for a cut-off in the BAT spectrum.

The public archival INTEGRAL data were analyzed to obtain a lightcurve. The IBIS lightcurve was constructed in the 18–60 keV energy band, and with individual points corresponding to about ~2000 s (INTEGRAL science window). We used the same methods as those used in the construction of the 4th IBIS Catalog (Bird et al. 2010). Individual science windows that had exposures of less than 500 s or where the source was at a large off-axis angle ( $>12^\circ$  where the instrument flux calibration is not optimized) were excluded from the lightcurve. This produced a lightcurve of 1295 observations, spanning MJD 52 704.16–54 604.20, with a cumulative exposure of ~2.6 Ms. The source is detected at  $5.3\sigma$  in the 18–60 keV band with a weighted mean flux of  $0.6 \pm 0.1 \text{ mCrab}$ ; consistent with the reported catalog flux of Bird et al. (2010). The

<sup>2</sup>  $\alpha = \Gamma - 1$ ; in the following, the subscript of spectral index  $\alpha$  denotes a single waveband with R for radio and X for X-ray; or a broad waveband with RX for a band extending from radio to X-ray, RO from radio to optical and OX from optical to X-ray.

<sup>3</sup> <http://heasarc.gsfc.nasa.gov/docs/swift/results/bs58mon/>



**Fig. 6.** The dot-dashed lines correspond to (from the innermost) 68%, 95% and 99% best fit confidence level contours for the HESS J1943+213. The positions of the possible counterparts are indicated: the partial circle in green/light grey is the (large) INTEGRAL error circle for IGR 19443+2117; the red/grey solid circle is radio beam size with NVSS J194356+211826 in center; the blue/grey dashed circle is the NRT position; the small magenta/grey solid circle is the ROSAT 90% position circle with 1RXH J194356.2+211824 in center; the triangle is used for NVSS position; the blue square is used for BAT position; the star is used for both CXOU J194356.2+211823 and 2MASS J19435624+2118233.

source is weak and its lightcurve is compatible with white noise. Fitting the constant to the lightcurve yields  $\chi^2/\text{d.o.f.} = 1.2$  for 1293 d.o.f.

Based on the facts that the ROSAT/HRI, *Chandra*, Swift/XRT and INTEGRAL/IBIS fluxes (within the error range of the cross-calibration constant) are constant and consistent with each other, that the ROSAT, *Chandra* and Swift sources are only 23'' away from the centroid of the H.E.S.S. source position, and that there are no other identified X-ray sources in the H.E.S.S. error circle, we conclude that these X-ray observatories detected the same object and that this object is a very likely counterpart to the H.E.S.S. source. Assuming this association is correct, in the following, we adopt the *Chandra* source location with its uncertainty for evaluation of potential counterparts in IR and radio bands.

### 3.2. IR, optical and UV counterpart

The 2MASS Catalog lists 19 infrared sources within the 90% H.E.S.S. error circle, however only one faint unidentified source 2MASS J19435624+2118233 is situated within the small *Chandra* error circle (Skrutskie et al. 2006; Landi et al. 2009; Tomsick et al. 2009). The angular separation between the *Chandra* position and the 2MASS object is 0.3''. The source is viewed across the Galactic Plane, therefore, its IR and optical fluxes are strongly affected by extinction. The magnitude lower limits in the *J* and *H* bands are 14.53 and 13.36, respectively, and the source magnitude in the *K* band is 13.98. The lower limit on optical and UV magnitudes of the source from Swift/UVOT observations are  $V > 20.3$  and  $UVM2 > 21.5$  (Landi et al. 2009). Based on maps of Galactic dust infrared emission of

Schlegel et al. (1998), the extinction coefficients in the direction of the source are  $A_V = 8.96$  mag,  $A_J = 2.49$  mag,  $A_H = 1.59$  mag,  $A_K = 1.02$  mag for a color excess  $E(B - V) = 2.81 \pm 0.25$ . The extinction corrected magnitudes are  $V_0 > 11.34$ ,  $J_0 > 12.04$ ,  $H_0 > 11.77$ ,  $K_0 = 12.96$ . The extinction correction of the UV magnitude, based on extinction curves of Cardelli et al. (1989), gives  $UVM2_0 > 26.04$  mag.

The measured color excess corresponds to a hydrogen column density  $N_H = (1.53 \pm 0.02) \times 10^{22} \text{ cm}^{-2}$  (Predehl & Schmitt 1995), which is higher than the  $N_H$  measured by the survey of Galactic neutral hydrogen,  $(8.4 \pm 1.3) \times 10^{21} \text{ cm}^{-2}$  (Kalberla et al. 2005). The  $N_H$  obtained from X-ray *Chandra* and Swift spectra are  $(1.89^{+0.25}_{-0.22}) \times 10^{22} \text{ cm}^{-2}$  and  $(1.37^{+0.12}_{-0.13}) \times 10^{22} \text{ cm}^{-2}$  respectively (Tomsick et al. 2009; Landi et al. 2009), thus consistent with the column density obtained from dust maps. Therefore, there is no conclusive evidence that the flux of the X-ray source is intrinsically absorbed (in spite of observed  $N_H$  being in excess of Galactic neutral hydrogen measurements). However, dust maps used to calculate extinction may not be precise at Galactic latitudes lower than  $5^\circ$  (Schlegel et al. 1998).

The potential IR counterpart to IGR J19443+2117 was observed at ESO by Masetti et al. (in prep.), during a campaign aimed at identifying INTEGRAL sources. A preliminary IR spectrum of 2MASS J19435624+2118233 obtained with SofI spectro-imager at NTT (Moorwood et al. 1998) is heavily reddened and does not show any readily apparent emission lines.

### 3.3. Radio counterpart

There is only one radio source in the NVSS Catalog within the H.E.S.S. error circle (NVSS J194356+211826; Condon et al. 1998). It is located 24.7'' away from the best fit H.E.S.S. source position and 3.5'' from the *Chandra* source position. Although the *Chandra* source is outside of the NVSS error circle ( $\sim 0.5''$ ), it is still well within the radio PSF of 45''. The radio source has been detected by various survey programs between 327 and 4850 MHz at positions which differ from NVSS by 1.7''–40'' (Taylor et al. 1996; Gregory et al. 1996; Reich et al. 1984; Furst et al. 1990; Becker et al. 1991; Gregory & Condon 1991; Griffith et al. 1991). The spectral index calculated for nine combined measurements is  $\alpha_R = 0.32$  (Vollmer 2009) and the source flux density measured by NVSS at 1400 MHz is 102.6 mJy (Condon et al. 1998). There is no evidence for flux variation over the 12-year time span of the collected data.

Between March and May 2010 the radio source was monitored by the Nançay Radio Telescope (NRT), a meridian transit telescope with a main spherical mirror of 300 m  $\times$  35 m (Theureau et al. 2007). Thirteen observations were performed at two frequencies 1419 and 2392 MHz. The data processing consisted of selecting the best drift scans, background subtraction, straightening of the baseline by polynomial fit and application of a median smoothing. The calibration was corrected by the observations of the radio sources 3C 123, 3C 196, 3C 317 and PKS 0237-23 with the same setup as the main target.

The average statistical error for all observations is  $\sigma_{\text{stat}} = 4$  mJy. The systematic errors  $\sigma_{\text{sys}} = 17$  mJy at low frequency and 10 mJy at high frequency, take into account the uncertainty in the background subtraction and baseline contamination by secondary beams. The former mostly affects the low frequency observations due to larger beam size (more background sources are found within the beam together with the source of interest). The 3-months lightcurve is constant within the range of systematic errors. The constant fit to the data yields  $\chi^2/\text{d.o.f.} < 1$  at both frequencies. The average fluxes at low and high frequency are

111 ± 21 mJy and 86 ± 14 mJy, respectively. The average spectral index is  $\alpha_R = 0.59 \pm 0.16$ . The low frequency flux is in agreement with NVSS but the spectral index is somewhat higher than the archival observations.

### 3.4. Fermi/LAT data

The primary instrument onboard *Fermi*, the Large Area Telescope (LAT) is an electron-positron pair production telescope, featuring solid state silicon trackers and caesium iodide calorimeters, sensitive to high-energy (HE) gamma-ray photons from ~20 MeV to >300 GeV (Atwood et al. 2009). The data were nominally taken in survey mode; the observatory is rocked north and south on alternate orbits so that every part of the sky is observed for ~30 min every 3 h.

The LAT analysis dataset spanned MJD 54 682.7–55 319.4, corresponding to a period of ~21 months. The data were reduced and analyzed using the *Fermi* Science Tools v9r15 package. The standard onboard filtering, event reconstruction, and classification were applied to the data (Atwood et al. 2009), and for this analysis the high-quality (“diffuse”) photon event class is used. Throughout the analysis, the “Pass 6 v3 Diffuse” (P6\_V3\_DIFFUSE) instrument response functions are applied.

Events between 0.1–100 GeV from a 10 degree region of interest centred on the *Chandra* X-ray counterpart to HESS J1943+213 were extracted and a maximum likelihood analysis was performed on the data using the `gtlike` tool. A model was constructed containing all the point sources within 15° that are listed in the *Fermi*/LAT 1FGL Catalog (Abdo et al. 2010a); the point sources were modelled by a single power-law function and their parameters were frozen to those reported in the 1FGL Catalog unless they were within 10° of HESS J1943+213. Also incorporated were the currently recommended models for the Galactic diffuse emission (`gll_iem_v02.fit`) and isotropic backgrounds; the Galactic diffuse emission is fit with a power-law scaling. As HESS J1943+213 does not appear in the 1FGL Catalog an additional source was inserted into the model at the *Chandra* source location assuming a power-law spectrum. The likelihood analysis reveals that no significant source is detected at the *Chandra* source location i.e. at the best estimated position of the potential HESS J1943+213 counterpart. The 95% 0.1–100 GeV flux upper limits calculated using a Bayesian approach (Helene 1991) are  $<1 \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$  for an assumed spectral index  $\Gamma = 1.5$  and  $<6 \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$  for an assumed spectral index  $\Gamma = 2.0$ . It should be noted that this is a highly complex, crowded region in which there is a large contribution from the Galactic diffuse emission.

Both upper limits are shown in Fig. 7, where the limit for  $\Gamma = 1.5$  is plotted in dark brown, and  $\Gamma = 2.0$  in light brown. The comparison of the *Fermi*/LAT upper limit with the flux measured by H.E.S.S., implies a break in the HE/VHE gamma-ray spectrum located between 100 GeV and 500 GeV (or even higher if the spectrum is attenuated by the extragalactic background light), and a hard HE spectrum with  $\Gamma < 2.0$ .

## 4. Discussion

The positional coincidence between the new VHE gamma-ray source HESS J1943+213 and a radio, IR and X-ray source together with the lack of any other obvious candidate counterpart strongly suggests an association. Figure 7 shows the spectral energy distribution (SED) of HESS J1943+213 together with the

radio, 2MASS, *Chandra*, Swift, INTEGRAL, and *Fermi* data. The data are not simultaneous but no significant variability was found within any energy range, therefore, we proceed with a contemporaneous discussion of the entire dataset from radio to VHE gamma-rays.

Based on the X-ray, IR and radio properties of IGR J19443+2117, Landi et al. (2009) and Tomsick et al. (2009) argued briefly in favour of its AGN nature. The detection in VHE gamma-rays helps place tighter constraints on the source nature. Because of its point-like appearance in VHE gamma-rays, the source is most likely an AGN, a gamma-ray binary or a PWN (see Sect. 1). All options are discussed below in light of the multiwavelength properties presented in Sect. 3.

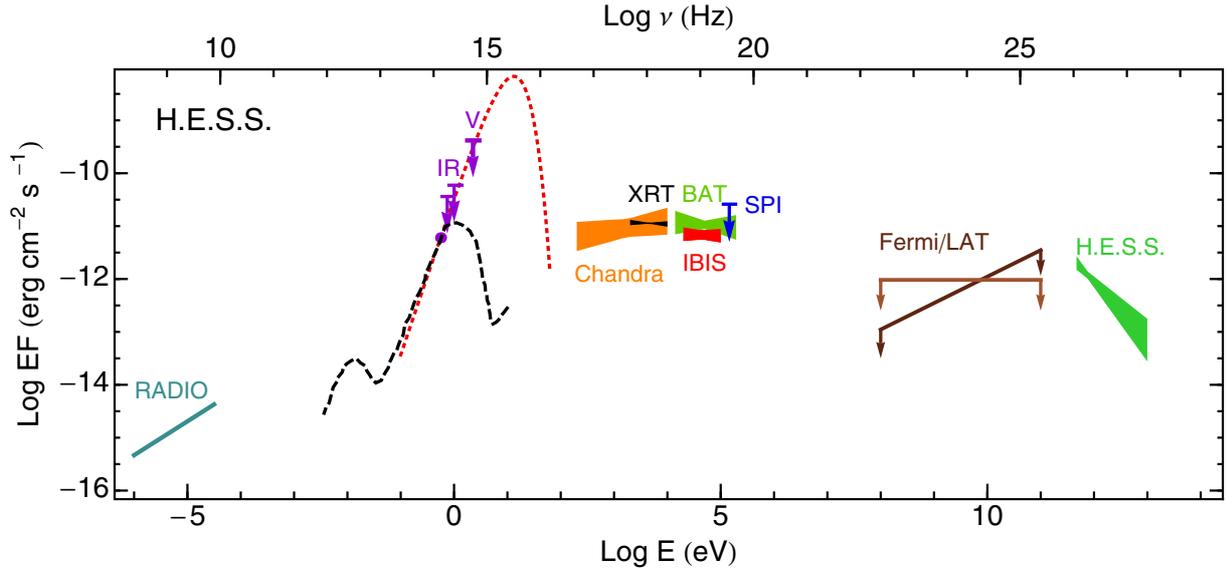
### 4.1. Gamma-ray binary

The point-like appearance of HESS J1943+213 in VHE gamma-rays and its location in the Galactic Plane may suggest a gamma-ray binary. With a handful of known systems, knowledge of these sources as a class is still limited. Known gamma-ray binaries are all radio sources with spectral indices at GHz frequencies  $\alpha_R \sim 0-0.5$  (Johnston et al. 1999; Ribó et al. 1999; Massi & Kaufman Bernadó 2009). Their X-ray spectra are hard ( $\alpha_X \sim 0.4-1.0$ ) with no cutoffs up to MeV energies (e.g., Takahashi et al. 2009; Smith et al. 2009; Uchiyama et al. 2009). The softest VHE spectra belong to LS I+61 303 ( $\Gamma = 2.7 \pm 0.5_{\text{stat}} \pm 0.2_{\text{sys}}$ , Albert et al. 2008) or PSR B1259-63 ( $\Gamma = 2.8 \pm 0.2_{\text{stat}} \pm 0.2_{\text{sys}}$ , Aharonian et al. 2009). In LS 5039 the photon index changes with the orbital phase between  $1.85 \pm 0.25$  and  $3.09 \pm 0.47$  (Aharonian et al. 2006a). The binaries LS 5039, LS I+61 303 and PSR B1259-63 have been detected by *Fermi*/LAT and the emission of the former two shows a cutoff at a few GeV, suggesting at least two components are required to explain their HE and VHE gamma-ray emission (Abdo et al. 2009a,c, 2010f). No HE counterpart has been found for HESS J0632+057.

The radio, X-ray and VHE spectral properties of HESS J1943+213 and its possible counterpart are compatible with what is known to date about gamma-ray binaries. The detection of orbital modulation would firmly associate HESS J1943+213 with a binary. However the lack of modulation in any of the spectral bands does not disqualify the HMXB scenario, since for example no modulation has been detected in HESS J0632+057.

All known gamma-ray binaries have a bright, massive star companion (O or Be spectral type) dominating the SED in the optical/IR band. Identifying the infrared source with a massive star would give credence to the identification as a binary. Furthermore, comparing the fluxes to those of the massive star in the gamma-ray binary LS 5039 (Fig. 3), the IR data are consistent with a massive O star only if its distance is >25 kpc, i.e. if it is located beyond the edge of our Galaxy (radius ~15 kpc) which is unlikely for a massive star. Similarly, comparing to the Be counterpart of PSR B1259-63 (the least luminous and the least hot among companions of known gamma-ray binaries) the lower limit on the distance is >22 kpc. A large distance would also imply a 1–10 keV X-ray luminosity of about  $10^{36} \text{ erg s}^{-1}$  whereas known gamma-ray binary luminosities are about  $10^{33}-10^{34} \text{ erg s}^{-1}$ .

Such a high X-ray luminosity can be explained if the source is accreting but, if the source is powered by a pulsar as in PSR B1259-63 and based on the observed correlations in pulsar wind nebulae (Mattana et al. 2009), this would require the spindown power of the pulsar to be higher than that of the Crab:  $\dot{E} \geq 10^{38} \text{ erg s}^{-1}$ . Finally, the possibility that the massive star is



**Fig. 7.** Spectral energy distribution of the sources discussed in Sect. 3. The details of the data and sources used in this plot are given in Table 1 and in Sect. 3.2. The dashed black line is the spectral template of elliptical galaxies from Silva et al. (1998) at a distance of 570 Mpc. The dotted (red) line is the blackbody spectrum of a massive star with parameters identical with the companion star in the gamma-ray binary LS 5039 at a distance of 25 kpc. The data in IR (2MASS) and visual (Swift/UVOT) bands were corrected for extinction. For *Fermi*/LAT we plot both upper limits for  $\Gamma = 1.5$  (dark brown) and  $\Gamma = 2.0$  (light brown).

strongly obscured by circumstellar material, as found in several high-mass X-ray binaries detected by INTEGRAL, appears unlikely as the X-ray derived  $N_{\text{H}}$  column density (see Sect. 3.2) is smaller than what is typically found in these systems (e.g. Rahoui et al. 2008).

Despite the general similarity of the radio to X-ray SED between HESS J1943+213 and known binaries, the lack of a plausible massive stellar counterpart in the optical/IR is the main argument which disfavors the binary hypothesis. The secondary argument against the binary hypothesis is the lack of orbital modulation in any of the bands.

#### 4.2. Pulsar wind nebula

When the wind ejected from a rotation-powered pulsar is confined by the pressure of the surrounding medium (supernova remnant or compressed interstellar gas), a PWN is created. PWNe constitute an abundant class of VHE Galactic sources detected by Cherenkov telescopes. Among them a number are unresolved (e.g., Aharonian et al. 2005a; Djannati-Ataï et al. 2008; Acciari et al. 2010), including one of the most powerful, the Crab Nebula (Aharonian et al. 2006b).

In general, the SED of PWNe consists of a broad synchrotron component extending up to high X-ray energies ( $\sim$ MeV) and an inverse Compton component at HE and VHE energies (Gaensler & Slane 2006). If HESS J1943+213 is a PWN, its gamma-ray (1–30 TeV) to X-ray (2–10 keV) flux ratio of 0.04 implies that the nebula is very young  $\sim 10^3$  yr, and that it contains a pulsar with a spindown power  $\dot{E} \sim 10^{38}$  erg  $s^{-1}$  (Mattana et al. 2009). The high  $\dot{E}$  is also supported by the X-ray photon index ( $\Gamma_{\text{X}} \approx 2$ ) (Gotthelf 2003).

The estimated PWN age and  $\dot{E}$  is similar to the Crab Nebula. Assuming that the luminosity of HESS J1943+213 is comparable to the Crab Nebula, we are able to estimate an approximate upper limit on the source distance. The low VHE flux of HESS J1943+213 being only 1.5% of the Crab Nebula (at 2 kpc; Trimble 1973) flux, implies a distance of  $\lesssim 16$  kpc. A comparison

with the remaining unresolved PWNe, which are also young but less luminous (10–15% of the Crab Nebula luminosity), would place HESS J1943+213 at a moderate distance of about 6 kpc.

The young age of HESS J1943+213 in the PWN scenario is consistent with it being unresolved for H.E.S.S.. However, all the unresolved VHE PWNe are extended in X-rays (Weisskopf et al. 2000; Ng et al. 2008; Matheson & Safi-Harb 2010; Temim et al. 2010), whereas HESS J1943+213 is point-like as observed by *Chandra* (Tomsick et al. 2009). Two unidentified H.E.S.S. sources, most likely associated with bright X-ray pulsars (Aharonian et al. 2006c) are embedded within very faint diffuse X-ray PWNe, these are HESS J1837-069 (AX J1838.0-0655, Gotthelf & Halpern 2008) and HESS J1616-508 (PSR J1617-5055, Kargaltsev et al. 2009). A 4.8 ks *Chandra* observation would miss such emission which could explain the point-like appearance of HESS J1943+213 in X-rays. However, both of the unidentified H.E.S.S. sources are extended in VHE gamma-rays and both have gamma-ray to X-ray flux ratios greater than unity, much higher than in case of HESS J1943+213.

The hard HE spectrum implied by the upper limits in Fig. 7 (see Sect. 3.4), would be in agreement with the HE spectrum of the Crab Nebula, where an IC component has  $\Gamma = 1.64$  above 1 GeV (Abdo et al. 2010e). Two other PWN detected by *Fermi*/LAT (Vela and HESS J1640-465) have soft spectra but these nebulae are older (more evolved) than the Crab, and are both extended in VHE gamma-rays. The VHE spectrum of HESS J1943+213 is significantly softer than all known VHE PWN. The softest PWN have photon indices of  $2.7 \pm 0.3$  and there are only a handful of them (Kargaltsev & Pavlov 2010). This fact together with the lack of obvious extended X-ray emission weaken the PWN hypothesis.

#### 4.3. Active galactic nucleus

Blazars are AGN where a relativistic jet pointing close to the line of sight produces Doppler-boosted emission (Blandford & Rees 1978; Urry & Padovani 1995). The spectral energy distribution

$\nu F_\nu$  of blazars is characterized by two non-thermal components: one peaking in the UV to X-ray range, commonly interpreted as synchrotron emission from highly relativistic electrons accelerated in the jet, and a second peaking in the gamma-ray regime being due to inverse Compton (IC) scattering of low-energy photons. Low-energy photons are the synchrotron radiation emitted either by the same population of electrons (synchrotron self-Compton, SSC), or ambient thermal photon (external Compton). The Doppler-boosted emission from the jet dominates the SED. When blazars are arranged according to their luminosity, the peak energy of both components is found to be anti-correlated with luminosity. High-frequency peak BL Lac (HBL) objects occupy the lowest end of the blazar luminosity sequence with the two spectral components peaking in X-rays (usually in the 0.1–1 keV range) and at about 100 GeV (Fossati et al. 1998). With the IC component peaking at HE, HBL objects are the best candidates for VHE sources, and indeed the class constitutes the majority of the current population of VHE AGN.

BL Lac objects are characterized in the optical/IR band by their elliptical host galaxy and have featureless optical spectra. Therefore, while the available preliminary NTT spectrum does not give a definite answer, the apparent lack of emission lines supports a classification of HESS J1943+213 as a BL Lac object.

If the 2MASS IR counterpart of HESS J1943+213 is an elliptical host galaxy, its flux can be compared with a template spectrum of the elliptical galaxies, (e.g. Silva et al. 1998). The measured flux in *K* band and the flux upper limits in the *H*, *J* and *V* bands, corrected for extinction, are consistent with an elliptical galaxy located at a distance  $>570$  Mpc ( $z \gtrsim 0.14$ ). This is a lower limit on the distance since non-thermal emission from the jet in a BL Lac object contributes a large fraction of the optical/IR flux. At such distances, absorption of VHE gamma rays on the extragalactic background light is expected to significantly soften the VHE spectrum. The observed H.E.S.S. spectrum is indeed quite soft (photon index of 3.2). The distance also sets the peak X-ray luminosity to be  $>10^{44}$  erg s $^{-1}$ . This is relatively high for a blazar but still consistent with spectra of BL Lac objects (Donato et al. 2001).

Broad band spectral indices are derived from the observations for an assumed host galaxy spectrum and compared to the typical values found in blazars. The indices are defined as  $\alpha_{RO} = -\log(F_O/F_R)/5.08$ ,  $\alpha_{OX} = -\log(F_X/F_O)/2.6$ ,  $\alpha_{RX} = -\log(F_X/F_R)/7.68$  where  $F_R$ ,  $F_O$  and  $F_X$  are the differential fluxes (per unit frequency) at 5 GHz, 5000 Å, and 1 keV respectively (Tananbaum et al. 1979). For HESS J1943+213 the indices are  $\alpha_{RO} = 0.36$ ,  $\alpha_{OX} = 0.96$  and  $\alpha_{RX} = 0.56$ . When compared to the values of  $\alpha_{RO}$  and  $\alpha_{OX}$  measured for a large group of blazars (e.g. Fig. 12 in Plotkin et al. 2010), HESS J1943+213 would be classified as a radio loud, X-ray strong HBL object. The radio loud classification of HESS J1943+213 is also supported by the value of  $\log R_X \equiv \log(\nu L_{6\text{ cm}}/L_{2-10\text{ keV}}) = -3.7$ , where the radio loud sources are those with  $\log R_X > -4.5$ , as defined by Terashima & Wilson (2003). The calculated value of  $\log R_X$  is also similar to the average value of  $-3.10$  calculated for a large sample of HBL objects, and is significantly different than the average  $-1.27$  for low-frequency peaked BL Lac objects and  $-0.95$  for flat spectrum radio quasars (Terashima & Wilson 2003).

Costamante & Ghisellini (2002) selected candidate VHE gamma-ray emitting BL Lac objects based on their fluxes in radio ( $F_R$ ) and X-ray ( $F_X$ ) bands measured at 5 GHz and 1 keV, respectively. BL Lac objects are expected to be detected at VHE energies if  $\log \nu_R F_R \gtrsim -14.8$  erg cm $^{-2}$  s $^{-1}$  and  $\log \nu_X F_X \gtrsim -11.4$  erg cm $^{-2}$  s $^{-1}$ . This target selection

criterion has proved quite successful (e.g. Aharonian et al. 2008; Abramowski et al. 2010). For HESS J1943+213 the fluxes are  $\log \nu_R F_R = -14.5$  and  $\log \nu_X F_X = -11.1$ , which places the source well within the group of VHE emitting BL Lac objects.

The X-ray spectrum of HESS J1943+213 is hard with no apparent cut-off up to 195 keV. In this respect HESS J1943+213 resembles extreme BL Lac objects, in which the synchrotron peak energy exceeds 10 keV (Costamante et al. 2001). The extreme status of HESS J1943+213 is also supported by  $\log(F_X/F_R) = -4.3$ , which is higher than  $-4.5$  as is the case for extreme HBL objects (Rector et al. 2003).

The extreme BL Lac objects are the least luminous blazars in the blazar sequence, and yet these are sources in which the dissipation of power carried by the jet is most efficient and particle acceleration mechanisms are close to their limits (Ghisellini 1999). The hard X-ray surveys (e.g. INTEGRAL, Swift) allow selection of extreme blazars (Ghisellini 2009). The group of known extreme HBL objects includes Mkn 501, 1ES 2344+514, PKS 0548-322 and 1ES 1426+428 (Costamante et al. 2001), all of which are VHE sources. The two spectral components of extreme BL Lac objects peak at energies  $\gtrsim 1$  keV and  $\gtrsim 100$  GeV for the synchrotron and IC component respectively. The HE spectrum in the 100 MeV–100 GeV domain is expected to be hard. The lack of an HE counterpart of HESS J1943+213 is unusual. Of all the VHE-detected AGN, only four (all HBL objects) have not been detected at GeV energies by *Fermi* (Abdo et al. 2010b). The upper limit by *Fermi* is quite strict but a similar situation has been observed e.g., in PKS 0548-322 (Aharonian et al. 2010).

The *Fermi*/LAT observations of TeV-selected AGN (Abdo et al. 2009b), show that their HE spectra are hard with photon indices  $\Gamma \lesssim 2.0$  where the extreme HBL objects belong to the hardest ones (i.e., 1ES 2344+514 with  $\Gamma = 1.57$  and 1ES 1426+428  $\Gamma = 1.49$ ). Taking into account the *Fermi*/LAT upper limit the predicted photon index  $\Gamma < 2$  of HESS J1943+213 fits well within the HBL hypothesis.

Blazars are a class of objects known for their variability on all timescales. Nevertheless, in the duty cycle of VHE-emitting AGN, there are periods without signs of significant flux variability. Accordingly, the apparent lack of flux variability in case of HESS J1943+213 does not contradict the hypothesis of it being an HBL object. For instance, HBL object PG 1553+113, detected in both HE and VHE gamma rays, has shown only weak variability in radio, optical and X-rays and no variability in gamma-rays despite regular monitoring of the sky above 200 MeV by *Fermi*/LAT (Abdo et al. 2010d). In addition, recent systematic studies of HE variability in blazars show that HBL objects are the least variable subgroup of blazars (Abdo et al. 2010c).

Based on the presented arguments, we conclude that the identification of HESS J1943+213 as an extreme HBL object is the most plausible hypothesis.

## 5. Conclusions

The H.E.S.S. array of Cherenkov telescopes has detected a new point-like VHE gamma-ray source located in the Galactic Plane. The source is not variable and has a soft spectrum. The VHE source is coincident with the unidentified INTEGRAL source IGR J19443+2117, which was also detected in soft X-rays, radio and IR bands. The positional coincidence and lack of any other obvious counterpart suggest that these are the same object. The available archival data on IGR J19443+2117 was combined with the new H.E.S.S., *Fermi* and radio observations.

This allowed us to discuss the source classification as a gamma-ray binary, PWN and AGN, the three known types of sources that appear point-like in VHE observations.

The classification as HMXB appears less likely given the absence of a massive star. A faint flux of the infrared counterpart attributed to a massive star and assuming a similarity of the source SED with the known binaries, implies too large a distance ( $>25$  kpc) and in consequence too high an X-ray luminosity. A secondary less constraining argument against the binary hypothesis is the absence of an orbital modulation.

The main argument against the PWN scenario is the soft VHE gamma-ray spectrum, which is unusual when compared with the spectra of a relatively large sample of TeV PWN. In addition, the unresolved X-ray emission as observed by *Chandra* weakens the PWN scenario.

There are no arguments against the AGN hypothesis. On the contrary, it is strongly supported by the radio, optical and X-ray properties being well within the parameter region (set by the broad band spectral indices and logarithmic flux ratios) in which high-frequency peaked BL Lac objects and TeV AGN are found. The infrared counterpart is compatible with an elliptical galaxy at  $z > 0.14$ , which would also explain the soft VHE spectrum as attenuation of the intrinsic spectrum on the extragalactic background light. The apparent lack of lines in the preliminary infrared spectrum (if confirmed) could be explained since jet continuum emission usually dominates in this energy range.

HBL objects comprise the majority of AGN detected until now in VHE gamma-rays. Statistically, assuming a uniform distribution on the sky, one out of 30 AGN can be expected to be situated at Galactic latitude  $|b| < 2^\circ$ . HESS J1943+213 seems to be one of the most extreme HBL objects detected, with a peak frequency well beyond 10 keV and a low X-ray to radio ratio.

The HBL classification would be secured by the identification of an elliptical galaxy in infrared images, the measurement of its redshift and the detection of correlated variability at radio, optical, X-ray or gamma-ray wavelengths. High resolution radio images from VLBI would also help to put further constraints on the source identity.

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## References

- Abdo, A. A., Ackermann, M., Ajello, M., et al. 2009a, *ApJ*, 701, L123 (Fermi LAT Collaboration)
- Abdo, A. A., Ackermann, M., Ajello, M., et al. 2009b, *ApJ*, 707, 1310 (Fermi LAT Collaboration)
- Abdo, A. A., Ackermann, M., Ajello, M., et al. 2009c, *ApJ*, 706, L56 (Fermi LAT Collaboration)
- Abdo, A. A., Ackermann, M., Ajello, M., et al. 2010a, *ApJS*, 188, 405 (Fermi LAT Collaboration)
- Abdo, A. A., Ackermann, M., Ajello, M., et al. 2010b, *ApJ*, 715, 429 (Fermi LAT Collaboration)
- Abdo, A. A., Ackermann, M., Ajello, M., et al. 2010c, *ApJ*, 722, 520 (Fermi/LAT Collaboration)
- Abdo, A. A., Ackermann, M., Ajello, M., et al. 2010d, *ApJ*, 708, 1310 (Fermi LAT Collaboration)
- Abdo, A. A., Ackermann, M., Ajello, M., et al. 2010e, *ApJ*, 708, 1254 (Fermi/LAT Collaboration)
- Abdo, A. A., Parent, D., Grove, J. E., et al. 2010f, *The Astronomer's Telegram*, 3085, 1
- Abramowski, A., Acero, F., Aharonian, F., et al. 2010, *A&A*, 516, A56 (H.E.S.S. Collaboration)
- Acciari, V. A., Aliu, E., Arlen, T., et al. 2010, *ApJ*, 719, L69 (VERITAS Collaboration)
- Aharonian, F., Akhperjanian, A. G., Aye, K., et al. 2004, *Astropart. Phys.*, 22, 109 (H.E.S.S. Collaboration)
- Aharonian, F., Akhperjanian, A. G., Barres de Almeida, U., et al. 2008, *A&A*, 481, L103 (H.E.S.S. Collaboration)
- Aharonian, F., Akhperjanian, A. G., Aye, K., et al. 2005a, *A&A*, 432, L25 (H.E.S.S. Collaboration)
- Aharonian, F., Akhperjanian, A. G., Aye, K., et al. 2005b, *A&A*, 442, 1 (H.E.S.S. Collaboration)
- Aharonian, F., Akhperjanian, A. G., Aye, K., et al. 2005c, *Science*, 307, 1938 (H.E.S.S. Collaboration)
- Aharonian, F., Akhperjanian, A. G., Bazer-Bachi, A. R., et al. 2006a, *A&A*, 460, 743 (H.E.S.S. Collaboration)
- Aharonian, F., Akhperjanian, A. G., Bazer-Bachi, A. R., et al. 2006b, *A&A*, 457, 899 (H.E.S.S. Collaboration)
- Aharonian, F., Akhperjanian, A. G., Bazer-Bachi, A. R., et al. 2006c, *ApJ*, 636, 777 (H.E.S.S. Collaboration)
- Aharonian, F., Akhperjanian, A. G., Bazer-Bachi, A. R., et al. 2006d, *A&A*, 460, 743 (H.E.S.S. Collaboration)
- Aharonian, F. A., Akhperjanian, A. G., Bazer-Bachi, A. R., et al. 2007, *A&A*, 469, L1 (H.E.S.S. Collaboration)
- Aharonian, F., Akhperjanian, A. G., Anton, G., et al. 2009, *A&A*, 507, 389 (H.E.S.S. Collaboration)
- Aharonian, F., Akhperjanian, A. G., Anton, G., et al. 2010, *A&A*, 521, A69 (H.E.S.S. Collaboration)
- Albert, J., Aliu, E., Anderhub, H., et al. 2006, *Science*, 312, 1771 (MAGIC Collaboration)
- Albert, J., Aliu, E., Anderhub, H., et al. 2008, *ApJ*, 684, 1351 (MAGIC Collaboration)
- Atwood, W. B., Abdo, A. A., Ackermann, M., et al. 2009, *ApJ*, 697, 1071 (Fermi LAT Collaboration)
- Baumgartner, W. H., Mushotzky, R., & Tueller, J. 2011, in *Amer. Astron. Soc. Meet. Abstr.*, 217, 125
- Becker, R. H., White, R. L., & Edwards, A. L. 1991, *ApJS*, 75, 1
- Bird, A. J., Malizia, A., Bazzano, A., et al. 2007, *ApJS*, 170, 175
- Bird, A. J., Bazzano, A., Bassani, L., et al. 2010, *ApJS*, 186, 1
- Blandford, R. D., & Rees, M. J. 1978, in *BL Lac Objects*, ed. A. M. Wolfe, 328
- Bouchet, L., Jourdain, E., Roques, J.-P., et al. 2008, *ApJ*, 679, 1315
- Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, 345, 245
- Chaves, R. C. G. 2009 [[arXiv:0907.0768](https://arxiv.org/abs/0907.0768)] (H.E.S.S. Collaboration)
- Condon, J. J., Cotton, W. D., Greisen, E. W., et al. 1998, *AJ*, 115, 1693
- Costamante, L., & Ghisellini, G. 2002, *A&A*, 384, 56
- Costamante, L., Ghisellini, G., Giommi, P., et al. 2001, *A&A*, 371, 512
- Cusumano, G., La Parola, V., Segreto, A., et al. 2010, *A&A*, 510, A48
- de Naurois, M., & Rolland, L. 2009, *Astropart. Phys.*, 32, 231
- Djannati-Ataï, A., de Jager, O. C., Terrier, R., et al. 2008, in *International Cosmic Ray Conference*, 2, 823 (H.E.S.S. Collaboration)
- Donato, D., Ghisellini, G., Tagliaferri, G., & Fossati, G. 2001, *A&A*, 375, 739
- Feldman, G. J., & Cousins, R. D. 1998, *Phys. Rev. D*, 57, 3873
- Fossati, G., Maraschi, L., Celotti, A., Comastri, A., & Ghisellini, G. 1998, *MNRAS*, 299, 433
- Furst, E., Reich, W., Reich, P., & Reif, K. 1990, *A&AS*, 85, 805
- Gaensler, B. M., & Slane, P. O. 2006, *ARA&A*, 44, 17
- Ghisellini, G. 1999, *Astropart. Phys.*, 11, 11
- Ghisellini, G. 2009, in *AIP Conf. Ser.* 1126, ed. J. Rodriguez, & P. Ferrando, 131

- Gotthelf, E. V. 2003, *ApJ*, 591, 361  
 Gotthelf, E. V., & Halpern, J. P. 2008, *ApJ*, 681, 515  
 Gregory, P. C., & Condon, J. J. 1991, *ApJS*, 75, 1011  
 Gregory, P. C., Scott, W. K., Douglas, K., & Condon, J. J. 1996, *ApJS*, 103, 427  
 Griffith, M., Heflin, M., Conner, S., Burke, B., & Langston, G. 1991, *ApJS*, 75, 801  
 Helene, O. 1991, *Nucl. Instr. Meth. Phys. Res. A*, 300, 132  
 Hinton, J. A., Skilton, J. L., Funk, S., et al. 2009, *ApJ*, 690, L101  
 Johnston, S., Manchester, R. N., McConnell, D., & Campbell-Wilson, D. 1999, *MNRAS*, 302, 277  
 Kalberla, P. M. W., Burton, W. B., Hartmann, D., et al. 2005, *A&A*, 440, 775  
 Kargaltsev, O., Pavlov, G. G., & Wong, J. A. 2009, *ApJ*, 690, 891  
 Kargaltsev, O., & Pavlov, G. G. 2010, in *AIP Conf. Ser.* 1248, ed. A. Comastri, L. Angelini, & M. Cappi, 25  
 Landi, R., Stephen, J. B., Masetti, N., et al. 2009, *A&A*, 493, 893  
 Li, T.-P., & Ma, Y.-Q. 1983, *ApJ*, 272, 317  
 Malizia, A., Landi, R., Bassani, L., et al. 2007, *ApJ*, 668, 81  
 Massi, M., & Kaufman Bernadó, M. 2009, *ApJ*, 702, 1179  
 Matheson, H., & Safi-Harb, S. 2010, *ApJ*, 724, 572  
 Mattana, F., Falanga, M., Götz, D., et al. 2009, *ApJ*, 694, 12  
 Moorwood, A. F. M., Cuby, J. G., & Lidman, C. 1998, *Messenger*, 91, 9  
 Ng, C., Slane, P. O., Gaensler, B. M., & Hughes, J. P. 2008, *ApJ*, 686, 508  
 Ohm, S., van Eldik, C., & Egberts, K. 2009, *Astropart. Phys.*, 31, 383  
 Plotkin, R. M., Anderson, S. F., Brandt, W. N., et al. 2010, *AJ*, 139, 390  
 Predehl, P., & Schmitt, J. H. M. M. 1995, *A&A*, 293, 889  
 Rahoui, F., Chaty, S., Lagage, P., & Pantin, E. 2008, *A&A*, 484, 801  
 Rector, T. A., Gabuzda, D. C., & Stocke, J. T. 2003, *AJ*, 125, 1060  
 Reich, W., Fuerst, E., Haslam, C. G. T., Steffen, P., & Reif, K. 1984, *A&AS*, 58, 197  
 Ribó, M., Reig, P., Martí, J., & Paredes, J. M. 1999, *A&A*, 347, 518  
 ROSAT Scientific Team 2000, *VizieR Online Data Catalog*, 9028  
 Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, *ApJ*, 500, 525  
 Silva, L., Granato, G. L., Bressan, A., & Danese, L. 1998, *ApJ*, 509, 103  
 Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, 131, 1163  
 Smith, A., Kaaret, P., Holder, J., et al. 2009, *ApJ*, 693, 1621  
 Takahashi, T., Kishishita, T., Uchiyama, Y., et al. 2009, *ApJ*, 697, 592  
 Tananbaum, H., Avni, Y., Branduardi, G., et al. 1979, *ApJ*, 234, L9  
 Taylor, A. R., Goss, W. M., Coleman, P. H., van Leeuwen, J., & Wallace, B. J. 1996, *ApJS*, 107, 239  
 Temim, T., Slane, P., Reynolds, S. P., Raymond, J. C., & Borkowski, K. J. 2010, *ApJ*, 710, 309  
 Terashima, Y., & Wilson, A. S. 2003, *ApJ*, 583, 145  
 Theureau, G., Hanski, M. O., Coudreau, N., Hallet, N., & Martin, J. 2007, *A&A*, 465, 71  
 Tomsick, J. A., Chaty, S., Rodriguez, J., Walter, R., & Kaaret, P. 2009, *ApJ*, 701, 811  
 Trimble, V. 1973, *PASP*, 85, 579  
 Uchiyama, Y., Tanaka, T., Takahashi, T., Mori, K., & Nakazawa, K. 2009, *ApJ*, 698, 911  
 Urry, C. M., & Padovani, P. 1995, *PASP*, 107, 803  
 Vollmer, B. 2009, *VizieR Online Data Catalog*, 8085, 0  
 Weisskopf, M. C., Hester, J. J., Tennant, A. F., et al. 2000, *ApJ*, 536, L81
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